Shattering and Coagulation of Dust Grains in Interstellar Turbulence

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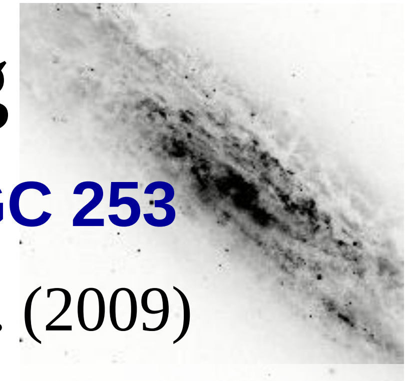
Outline

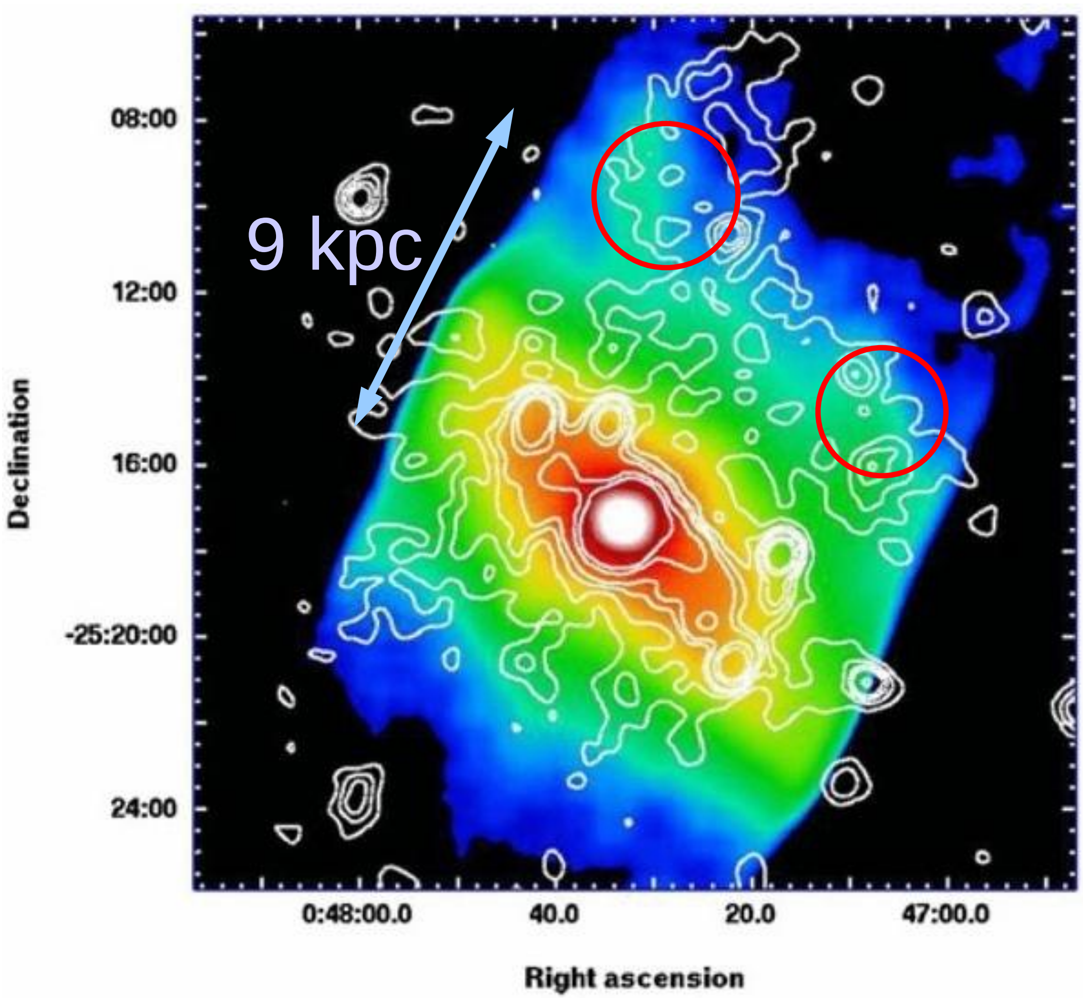
1. Dynamical Grain-Gas Coupling
2. Interstellar Turbulence
3. Formulation of Shattering and Coagulation
4. In a Cosmological Context

1. Dynamical Grain-Gas Coupling



NGC 253

Kaneda et al. (2009)



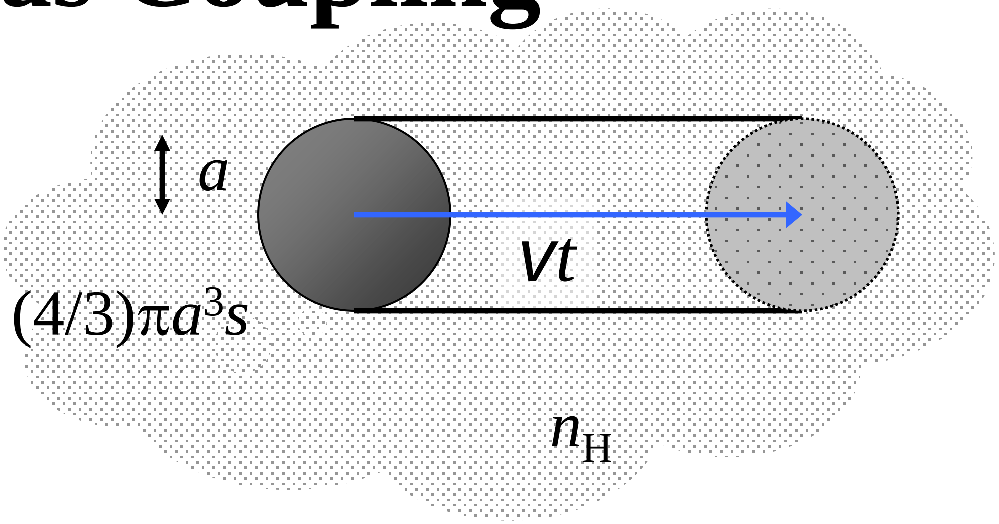
color: *AKARI* 90 μm
contour: *ROSAT* X-ray

The FIR extension coincides with the X-ray structure.

Dust is transported by gas ejection induced by stellar activity.

V_{escape} = 280 km/s
9 kpc / 280 km/s = 30 Myr

Grain-Gas Coupling



$m_{gr} = (4/3)\pi a^3 s$

Gas drag timescale t_d:

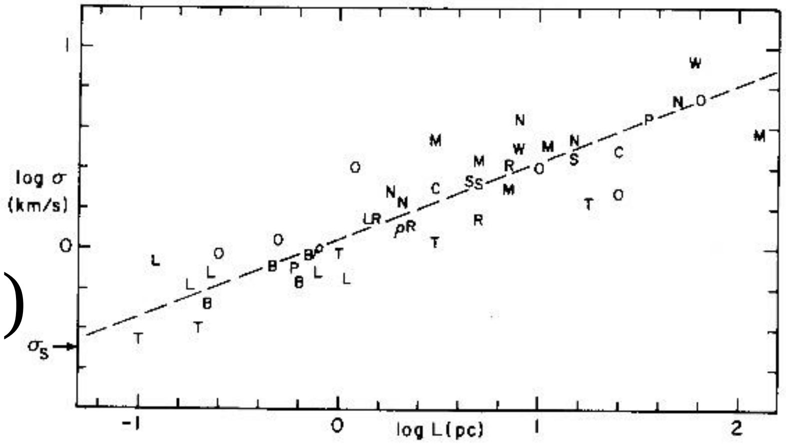
$(m_H v)(\pi a^2 v n_H) t_d \sim m_{gr} v$.

Grain motion is coupled with the gas motion on a scale l large enough:

$l \sim v t_d = (4/3) a s / (m_H n_H) \sim (10/n_H)(a/0.1\ \mu\text{m})$ pc

Large grains tend to be coupled with larger motions.

2. Interstellar Turbulence



Larson (1981)

ISM is turbulent (often supersonic) (e.g., McKee & Ostriker 2007).

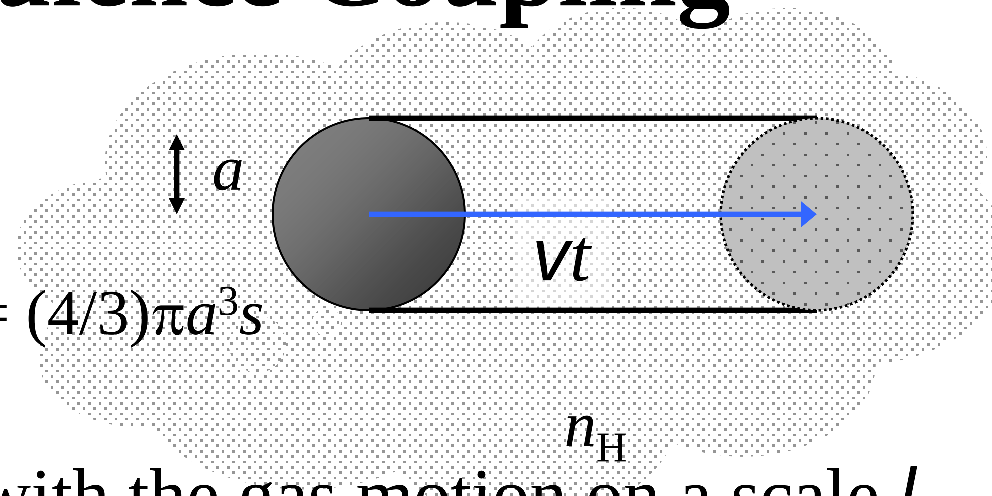
Implication for shattering (disruption):

c_s ~ 10 km/s in warm (~ 8000 K) medium → above the shattering threshold (~ a few km/s).

Implication for coagulation (sticking):

v_{turb} >> grain thermal speed. → If grain motion is coupled with turbulence, grain-grain collision occurs frequently (e.g., Ossenkopf 1993).

Grain-Turbulence Coupling



$m_{gr} = (4/3)\pi a^3 s$

Grain motion is coupled with the gas motion on a scale l large enough:

$l \sim vt_d = (4/3)as/(m_H n_H) \sim (10/n_H)(a/0.1\ \mu\text{m})$ pc

Large grains tend to be coupled with larger motions.

Kolmogorov turbulence: $v_{turb} \propto l^{1/3}$

Large grains tend to acquire larger velocities.

Gyroresonance

Magnetic fields ($B^2/8\pi \sim nkT$) + Grain charge

$\rightarrow$ MHD wave + gyro-motion of grains

Resonance between wave and gyro-motion (gyroresonance): $\omega - k_{//} v \cos\theta = n\omega_{\mathrm{gyro}}$

Large grains are further accelerated.

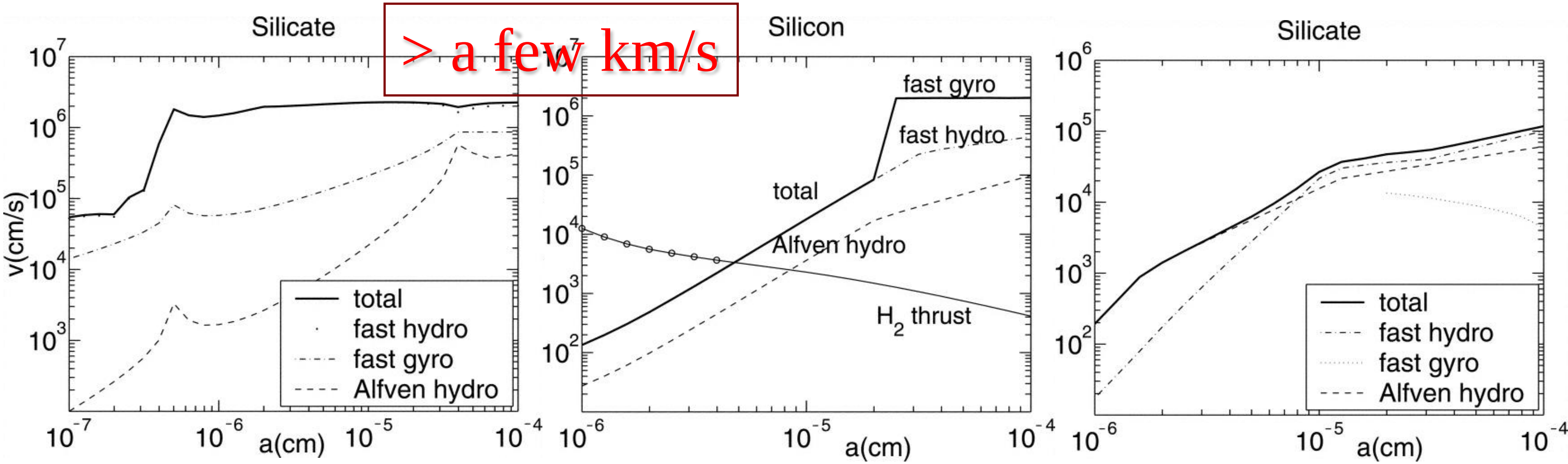
Grain Velocities

MHD turbulence model

hydro-drag, gyro-resonance

Yan, Lazarian, & Draine (2004)

> a few km/s

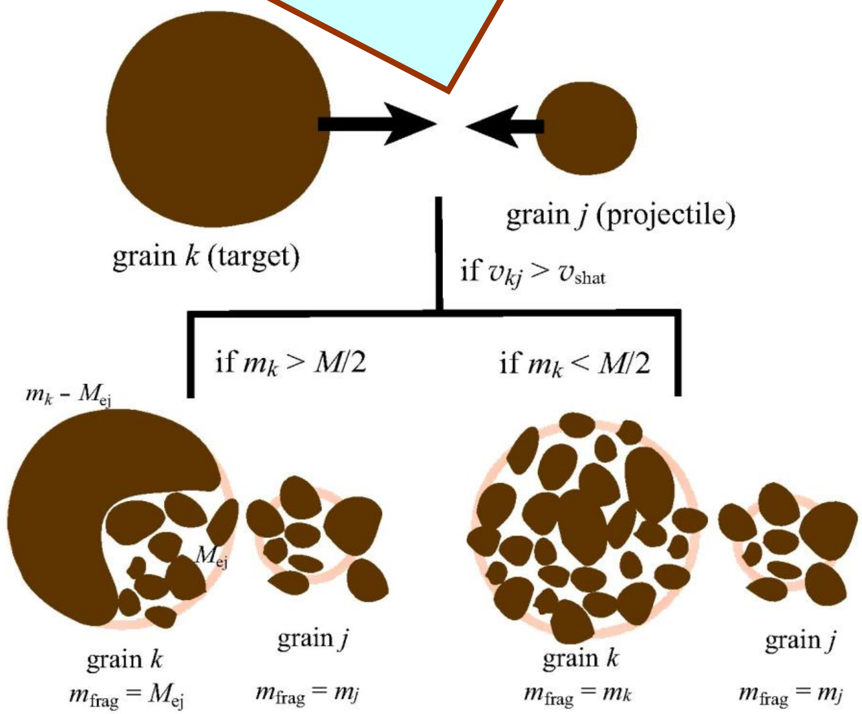


Warm ionized medium	Warm neutral medium	Dense cloud
T = 8000 K	T = 6000 K	T = 10 K
$n_H = 0.1$ cm^{-3}	$n_H = 0.3$ cm^{-3}	$n_H = 10^4$ cm^{-3}
B = 3.4 μG	B = 5.8 μG	B = 80 μG

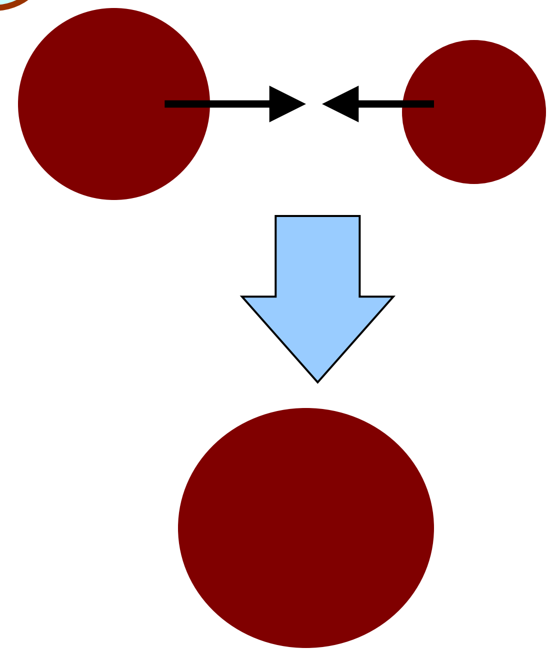
Sh… …gulation

Relative velocities can be excited by interstellar turbulence.

Shatte… …gulation



Shattering threshold:
2.7 km/s (silicate), 1.2 km/s (graphite)
(Jones et al. 1996)



coagulation rate = grain-grain collision rate
Threshold: ~ 10^3 cm/s

Grain Size Distribution and Extinction

Extinction (absorption+scattering)

$$\tau_{\lambda,i} = \int_0^\infty \pi a^2 Q_\lambda(a) N_{\rm dust}(a)\,da$$

$$\tau_\lambda = \sum_i \tau_{\lambda,i}$$

a: grain radius (<~ 0.1 μm)
$Q_\lambda(a)$ ~ 1 for λ <~ a
$Q_\lambda(a)$ << 1 ∝ a for λ >> a

i: grain species (silicate, graphite)

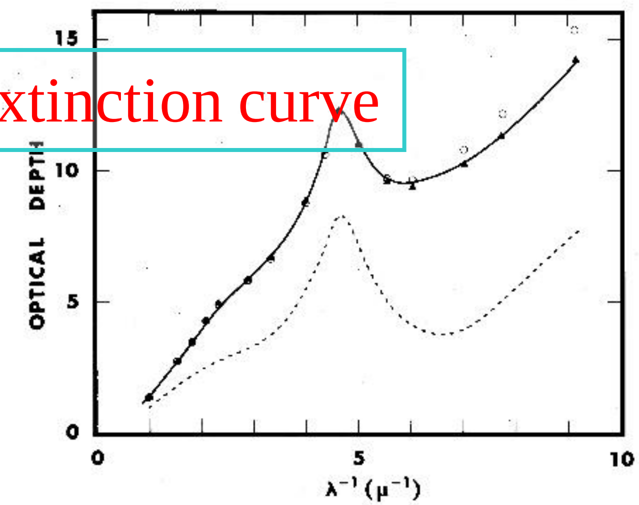


FIG. 4.—Optical depths, for column densities of 10^{22} H atoms cm^{-2}, versus inverse wavelength. *Solid line:* observed by OAO. *Triangles:* the extinction of (C + Ol) mixture of Fig. 2. *Dashed line:* the contribution of graphite to the extinction. *Dots:* a mixture of graphite and olivine, $n(a) \propto a^{-3.5}$, 0.005 μm < a < 0.25 μm, forced to fit at the maximum of the "bump" at 4.6 μm^{-1}.

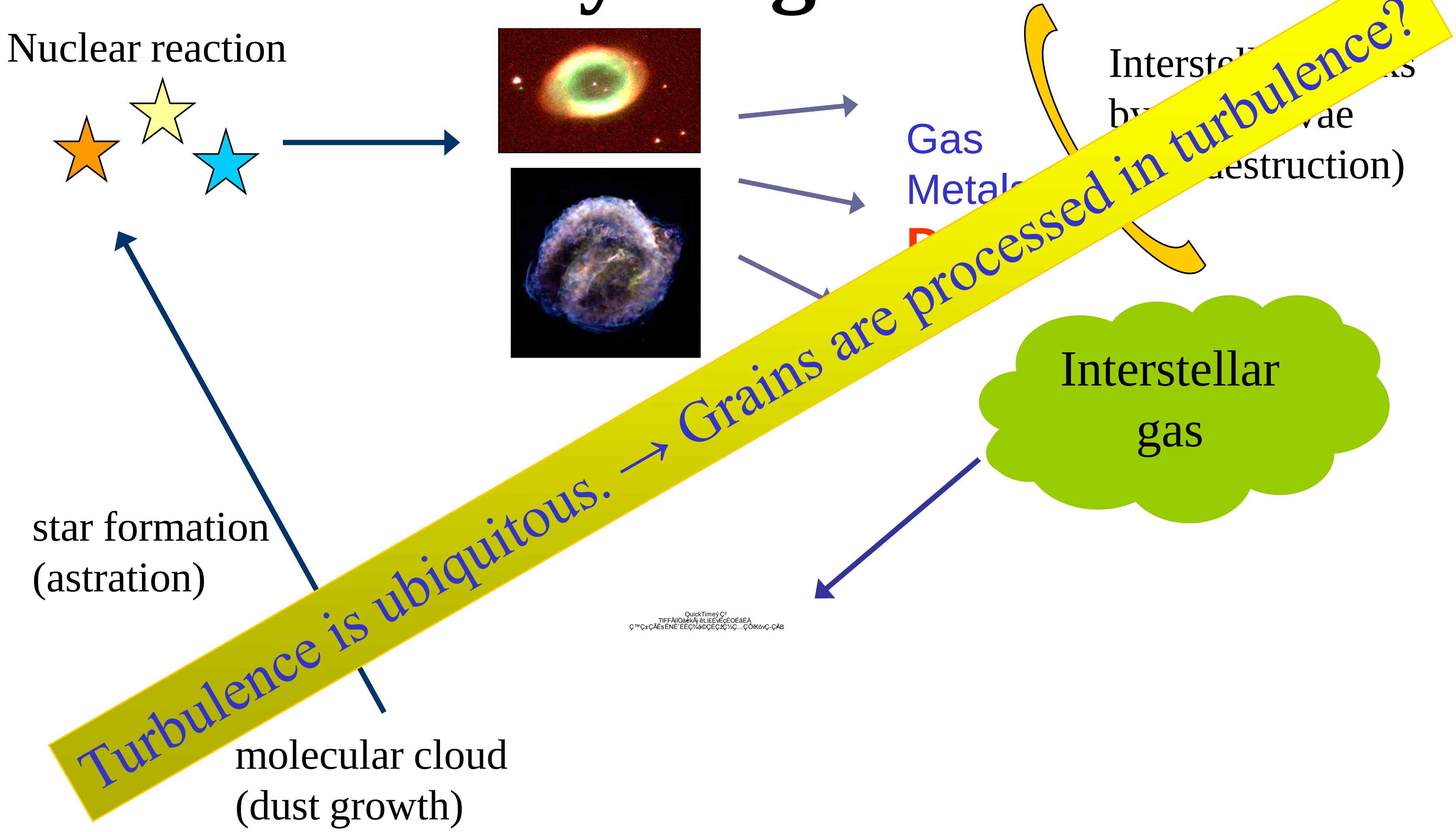
Mathis, Rumpl, & Nordsieck (1977)

Grain size distribution $N_{\rm dust}(a) \propto a^{-3.5}$
with 0.005 μm < a < 0.25 μm: MRN

What determines the grain size distribution?

- ☐ Source (supernova, AGB stars, etc.)
- ☐ **Shattering and coagulation?**

ISM Recycling and Dust



Specific Questions

A) Evolution of grain size distribution by shattering and coagulation under the grain motion induced by *ubiquitous* turbulence.

B) Do shattering and coagulation have a significant imprint in the extinction curve?

3. Formulation

Hirashita & Yan (2009)

Discrete size bins $a_0, ..., a_N$
The i-th bin contains grains of $\tilde{\rho}_i$ [g cm^{-3}].

Shattering

$$\left[\frac{\mathrm{d}\tilde{\rho}_i}{\mathrm{d}t}\right]_{\mathrm{shat}} = -m_i\tilde{\rho}_i\sum_{k=1}^{N}\alpha_{ki}\tilde{\rho}_k + \sum_{j=1}^{N}\sum_{k=1}^{N}\alpha_{kj}\tilde{\rho}_k\tilde{\rho}_j m_{\mathrm{shat}}^{kj}(i),$$

$$\alpha_{ki} = \begin{cases} \dfrac{\sigma_{ki}v_{ki}}{m_i m_k} & \text{if} \quad v_{ki} > v_{\mathrm{shat}}, \\ 0 & \text{otherwise,} \end{cases}$$

distribution of shattered fragments (power-law)

Coagulation

$$\left[\frac{\mathrm{d}\tilde{\rho}_i}{\mathrm{d}t}\right]_{\mathrm{coag}} = -m_i\tilde{\rho}_i\sum_{k=1}^{N}\alpha_{ki}\tilde{\rho}_k + \sum_{j=1}^{N}\sum_{k=1}^{N}\alpha_{kj}\tilde{\rho}_k\tilde{\rho}_j m_{\mathrm{coag}}^{kj}(i),$$

$$\alpha_{ki} = \begin{cases} \dfrac{\sigma_{ki}v_{ki}}{m_i m_k} & \text{if} \quad v_{ki} < v_{\mathrm{coag}}^{ki}, \\ 0 & \text{otherwise.} \end{cases}$$

$m_k + m_j$

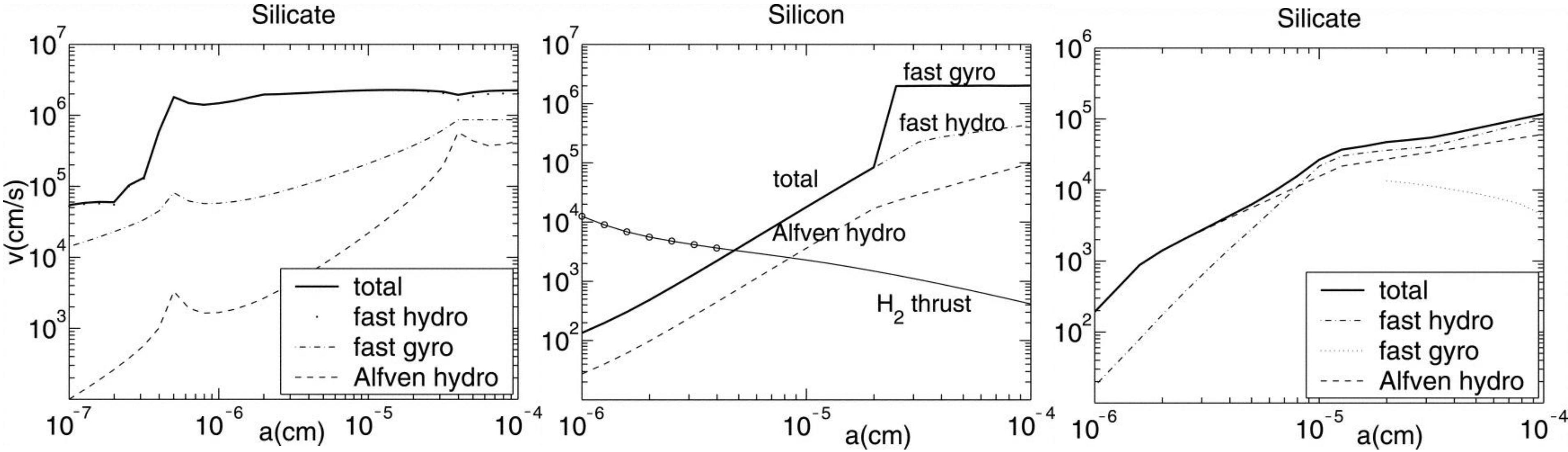
The grain velocities are adopted from Yan et al. (2004) (MHD turbulence).

Grain Velocities

MHD turbulence model

hydro-drag, gyro-resonance

Yan, Lazarian, & Draine (2004)



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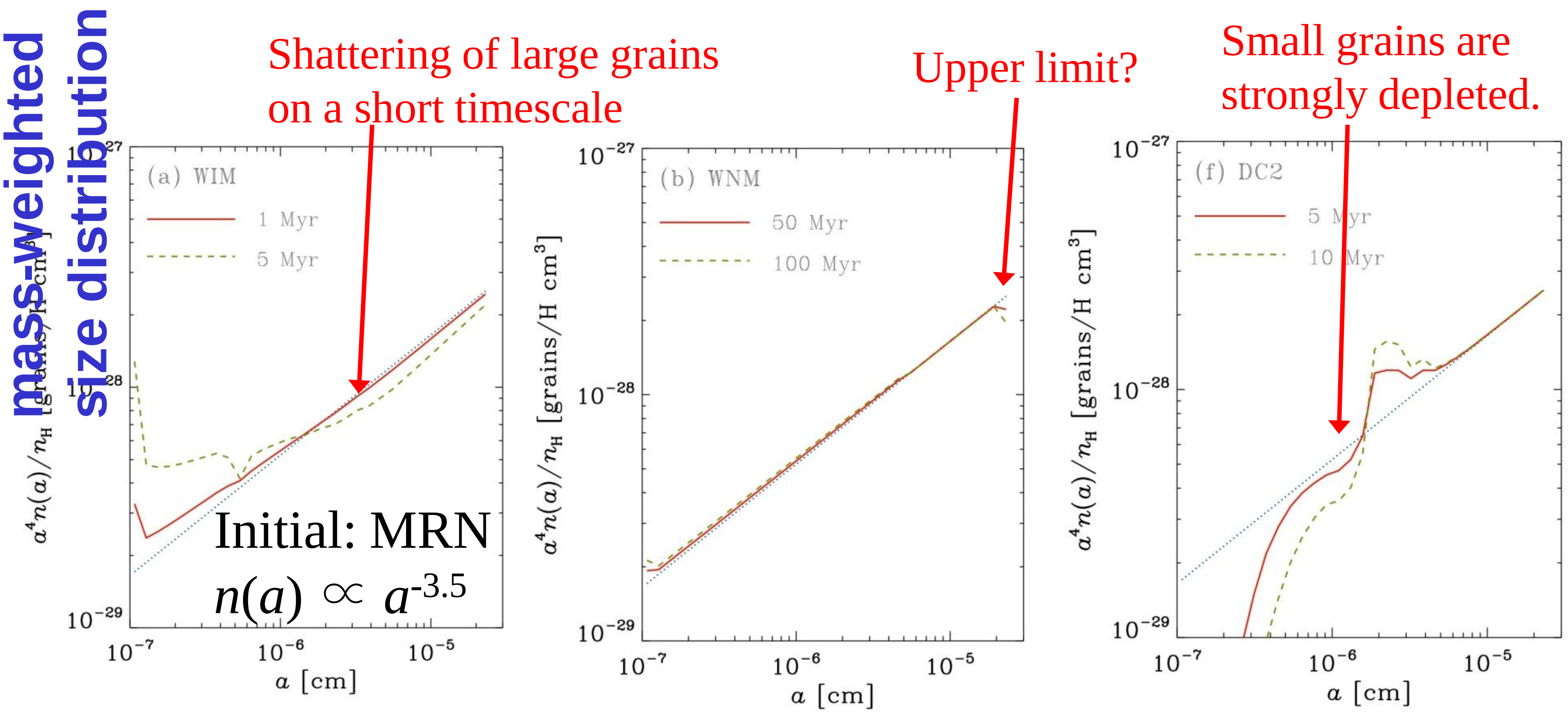
n_H = 0.1 cm^{-3}	n_H = 0.3 cm^{-3}	n_H = 10^4 cm^{-3}
B = 3.4 μG	B = 5.8 μG	B = 80 μG

Results



Warm ionized medium
T = 8000 K
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Warm neutral medium
T = 6000 K
n_H = 0.3 cm^{-3}
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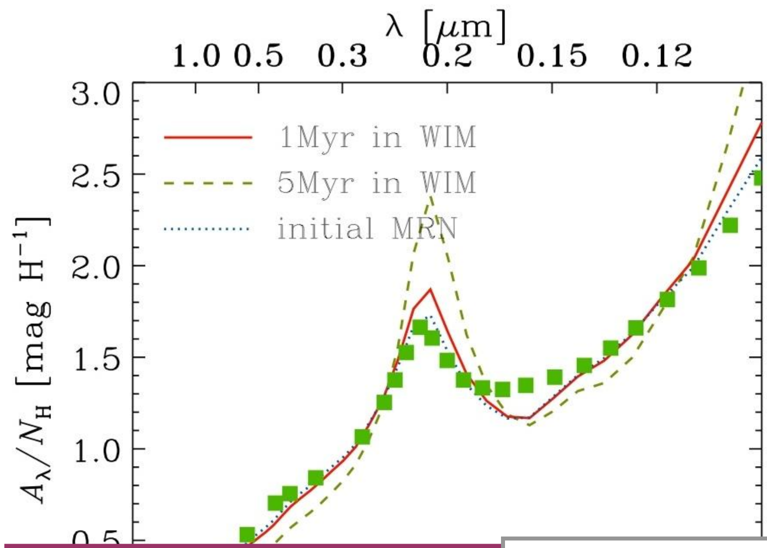
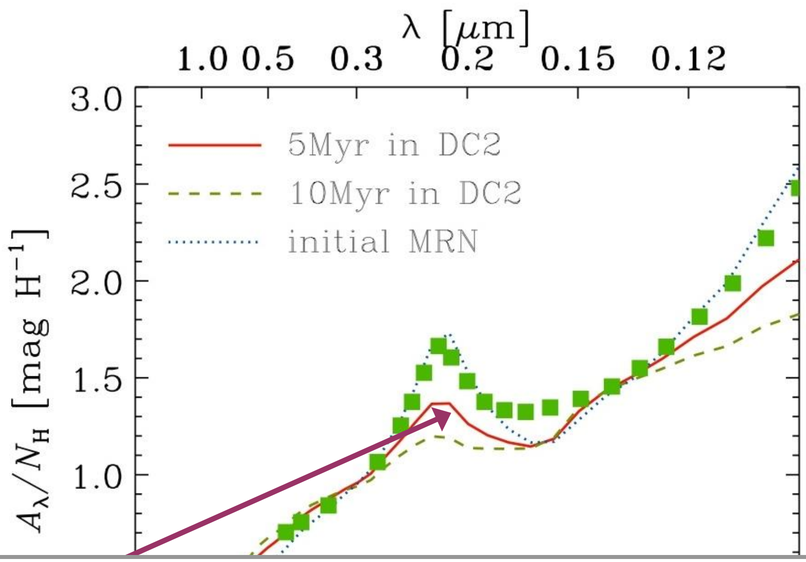
Dense cloud
T = 10 K
n_H = 10^4 cm^{-3}
B = 80 μG

Effects on the Extinction Curves





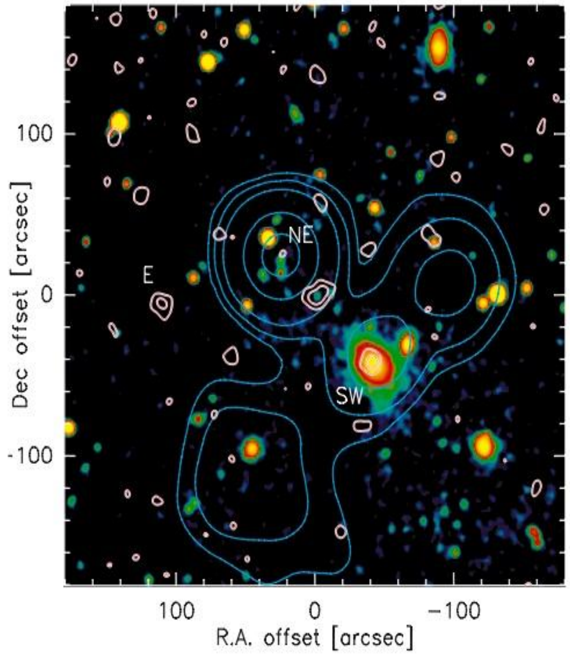
(1) The UV slope
the bump stren
sense.

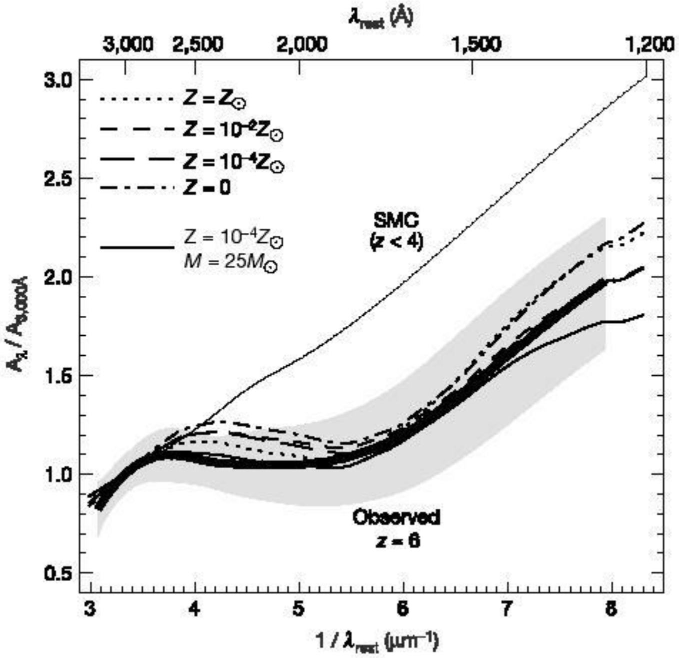
(2) The central pos
carbon bump is

(3) Small variation in IR
extinction curve.

Shattering and Coagulation in ISM can regulate the grain size distribution (and the extinction curve).

4. In a Cosmological Context





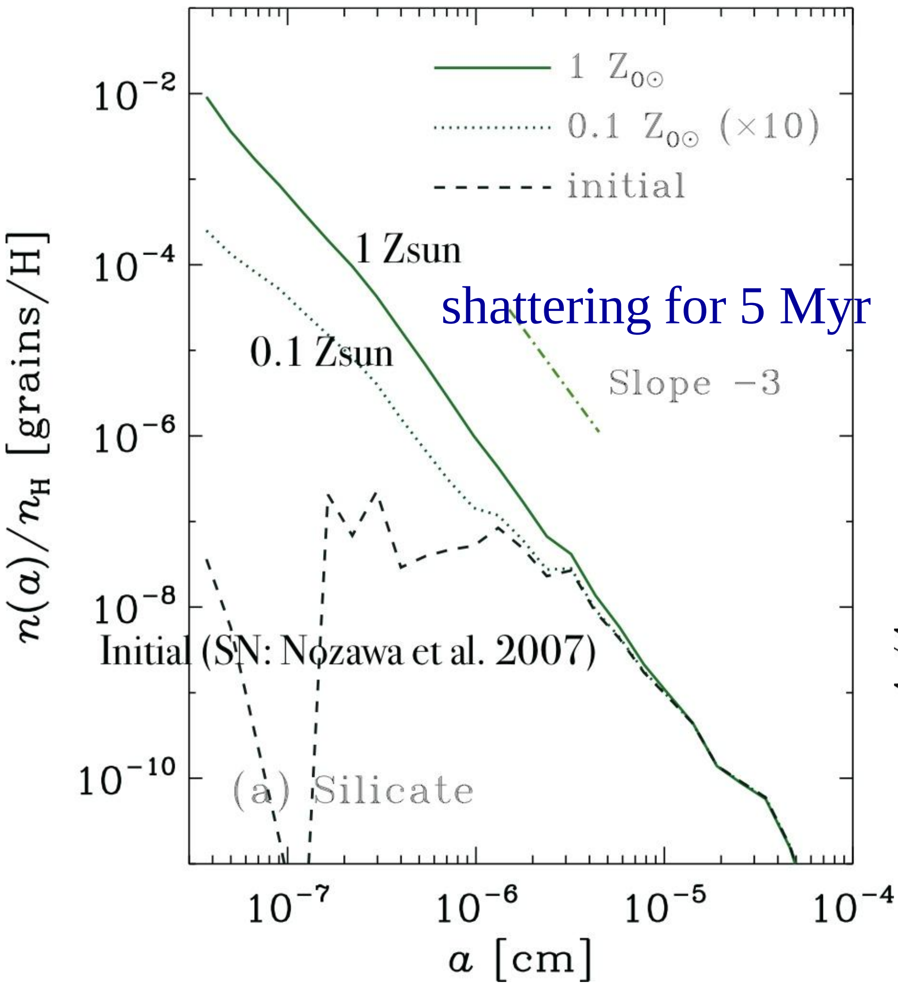
Dust already existed at
z ~ 6 (Bertoldi et al. 2003).

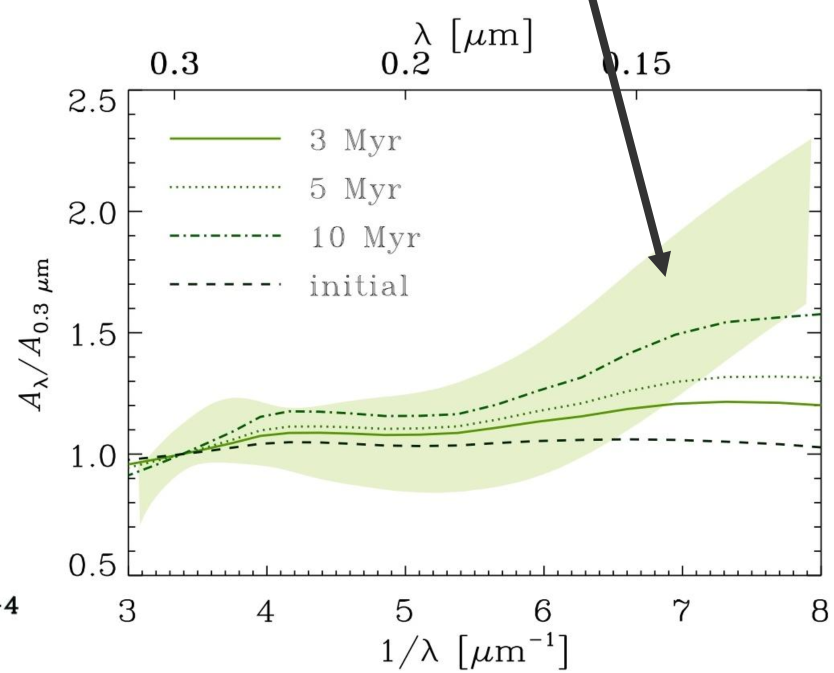
Extinction curve at *z* ~ 6
Maiolino et al. (2004)

Suggesting a rapid dust enrichment by supernovae

Shattering of SN Dust

Hirashita et al. (2010)
Hirashita (2010)





Small grain production by shattering contributes to the steepness of the UV extinction curve (**in solar metallicity**).

Scenario as a Summary

(1) The grain size distribution in the formation by supernovae (or AGB stars) is not processed by turbulence if the metallicity is $\ll 1/10\ Z_{\odot}$.

(2) After the metallicity enrichment, grain processing in ISM should be considered.

(3) In considering the origin of the grain size distribution at the present cosmic age, **interstellar processing by turbulence should be important** (occurs everywhere → easy to explain the universality).

Thank you.

3. Effects of Coagulation on SF

Hirashita & Omukai (2009)

(1) How about the denser regime?

(2) Importance of dust grains in star formation:

A) H_2 formation (H_2 is an efficient coolant for $Z < 0.01$ $Z_\odot$) ⇒ The grain surface S is important.

B) Dust cooling ⇒ The grain opacity κ_P is important.

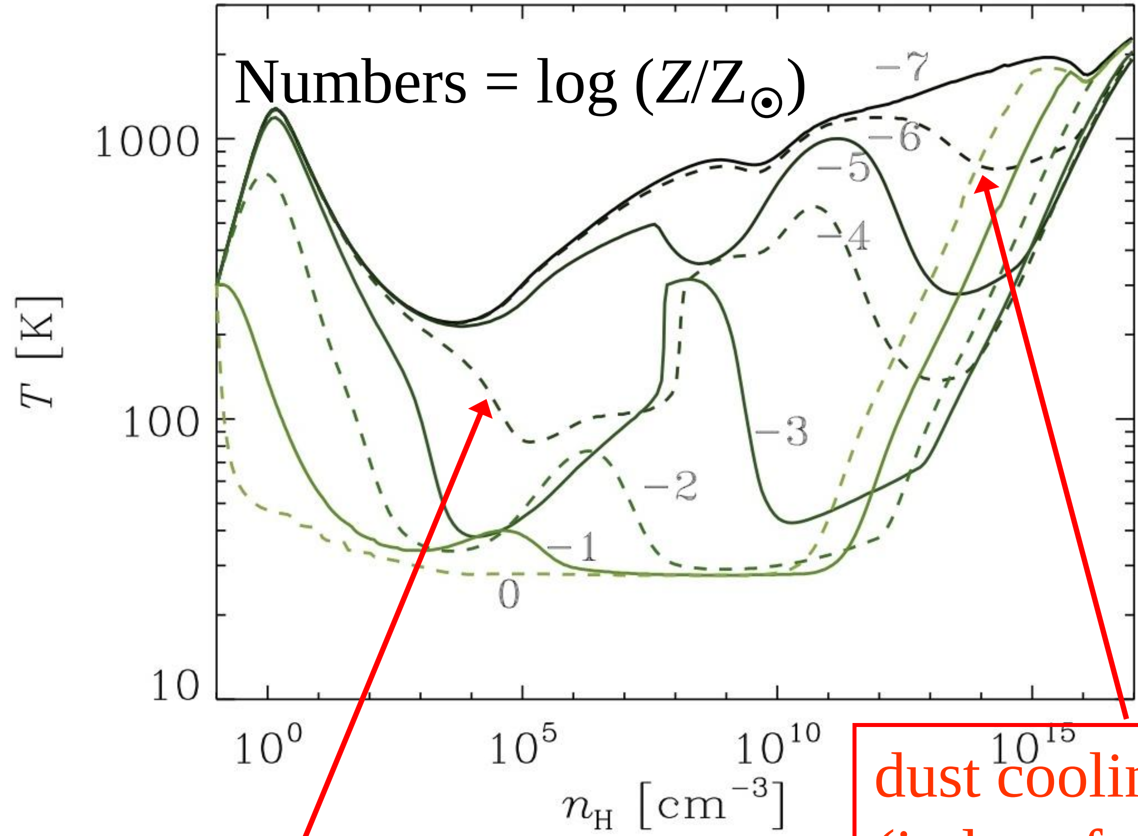
↓

We calculate the variation of S and κ_P in star-forming (collapsing) clouds.

Grain motion is assumed to be thermal.

Gas Evolution in Collapsing Clouds

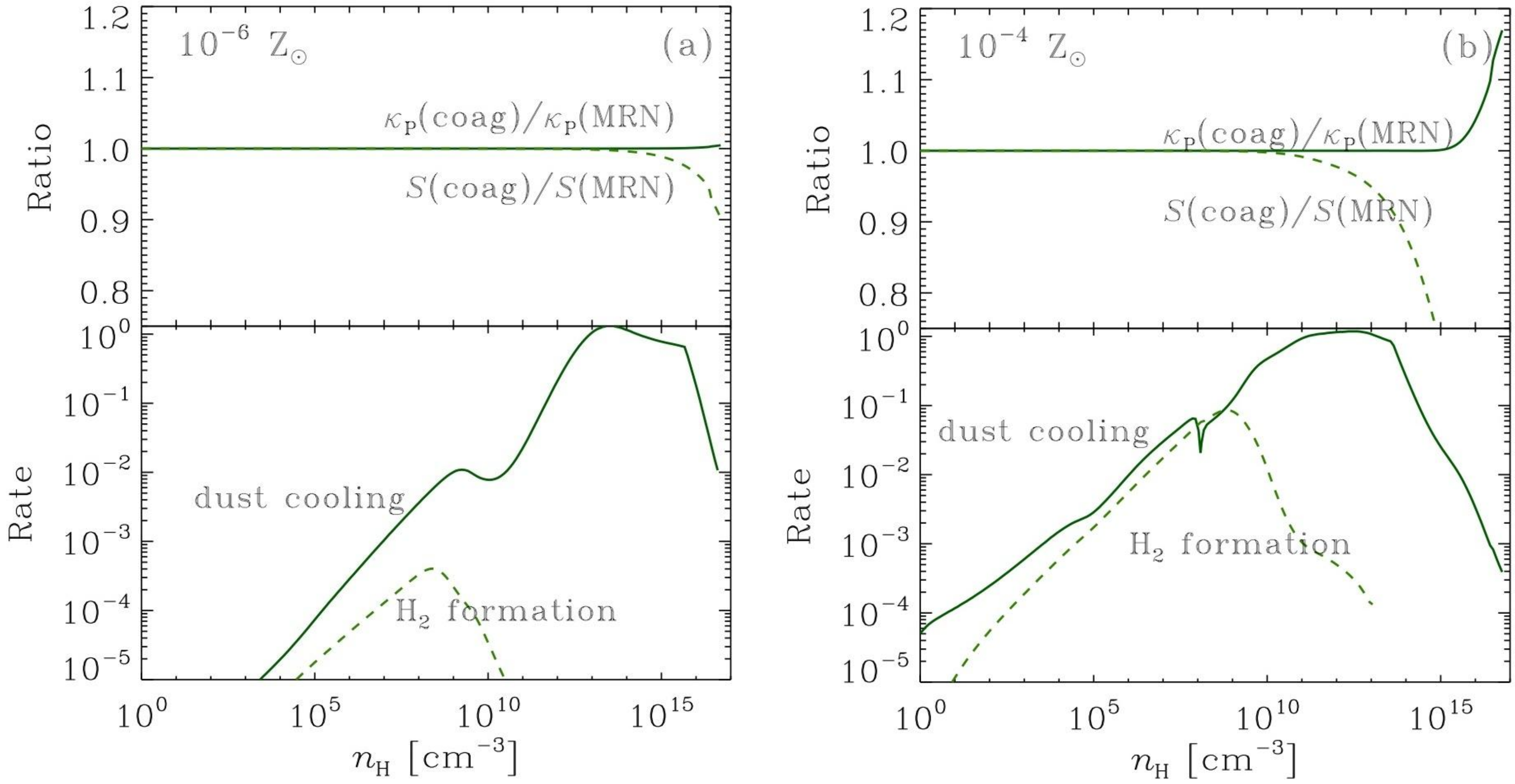
Omukai et al. (2005)



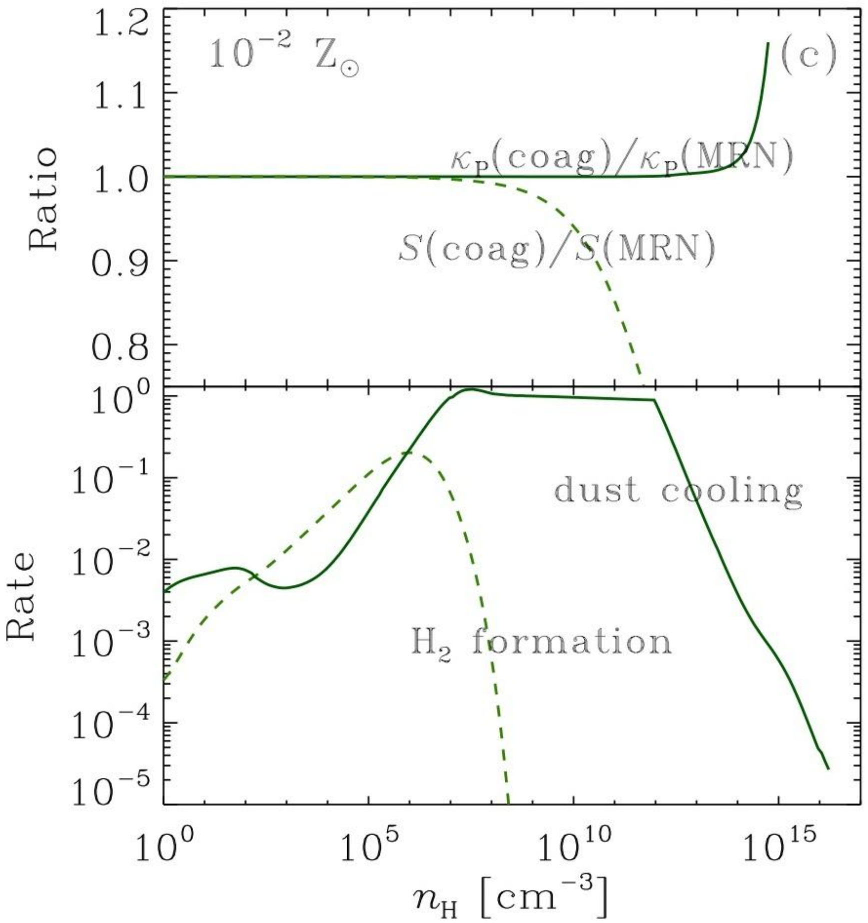
H_2 formation on grain surface: important coolant for log $(Z/Z_\odot) < -2$

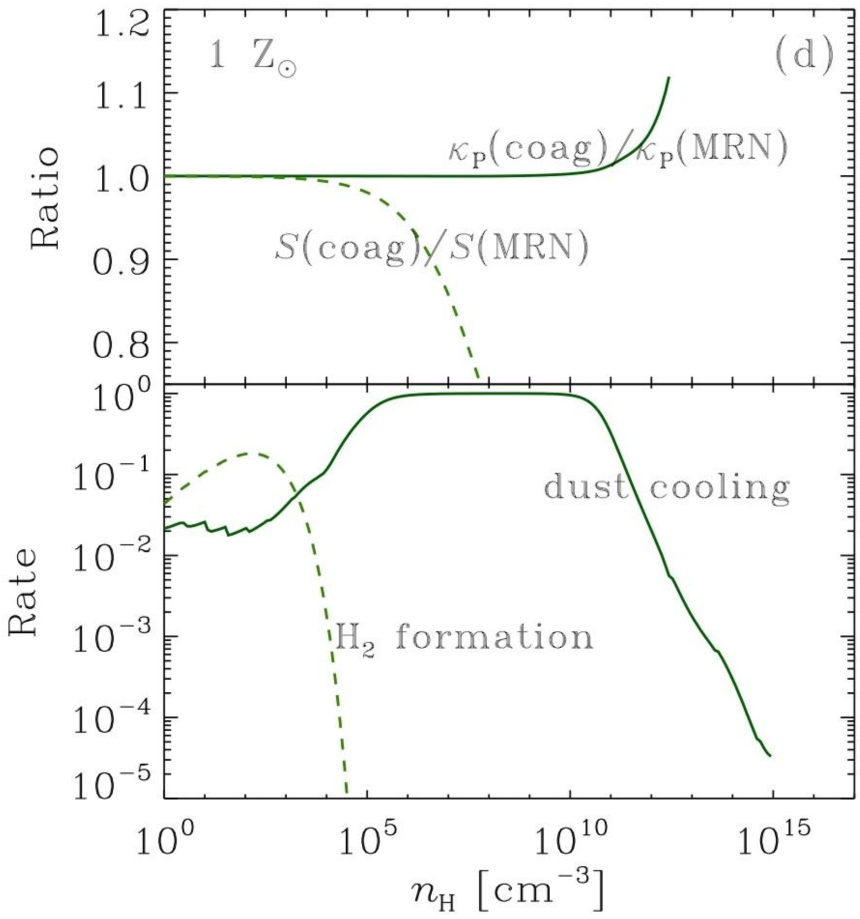
dust cooling (induce fragmentation)
Omukai et al. (2005)
Schneider et al. (2004)

Change of Grain Surface and Opacity by Coagulation



Change of Grain Surface and Opacity by Coagulation





Physical Considerations

☆ Grain surface is dominated by small grains. → Once the smallest grains are affected by coagulation, S begins to decrease (however, H_2 formation occurs faster).

- $t_{\rm ff} > t_{\rm coag} \Leftrightarrow n_{\rm H} > 10^7 (Z/Z_\odot)^{-2} (T/30\ {\rm K})^{-1}\ {\rm cm}^{-3}$

☆ Opacity ($\kappa_{\rm P} \propto \pi a^2 Q_\lambda \propto a^3$) is only a function of mass as long as $a << \lambda$. ⇒ $\kappa_{\rm P}$ does not change even if coagulation proceeds.

Coagulation has no effect on the thermal evolution in protostellar collapse.

1. Dust Grains in Galaxies